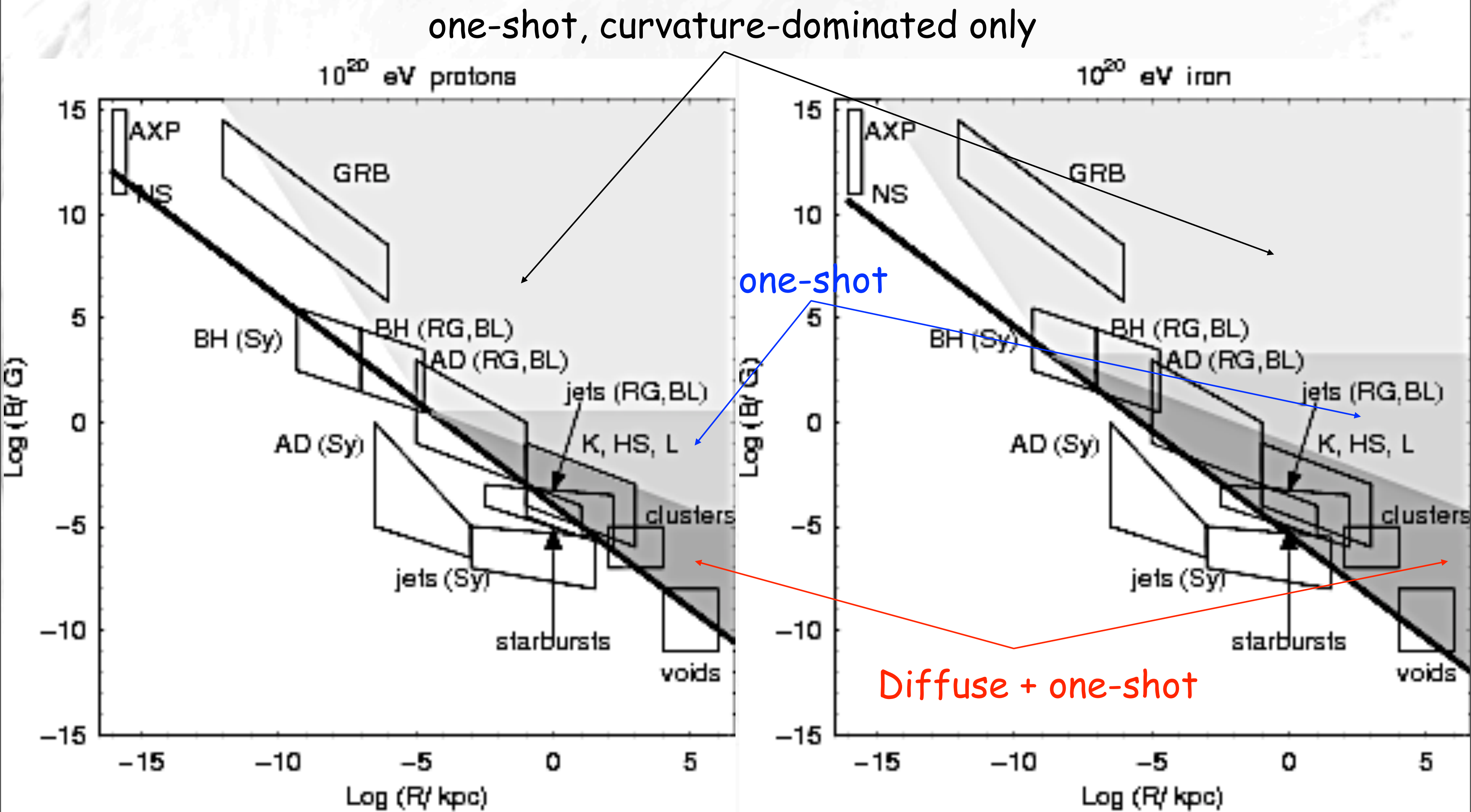
Hillas plot with energy losses

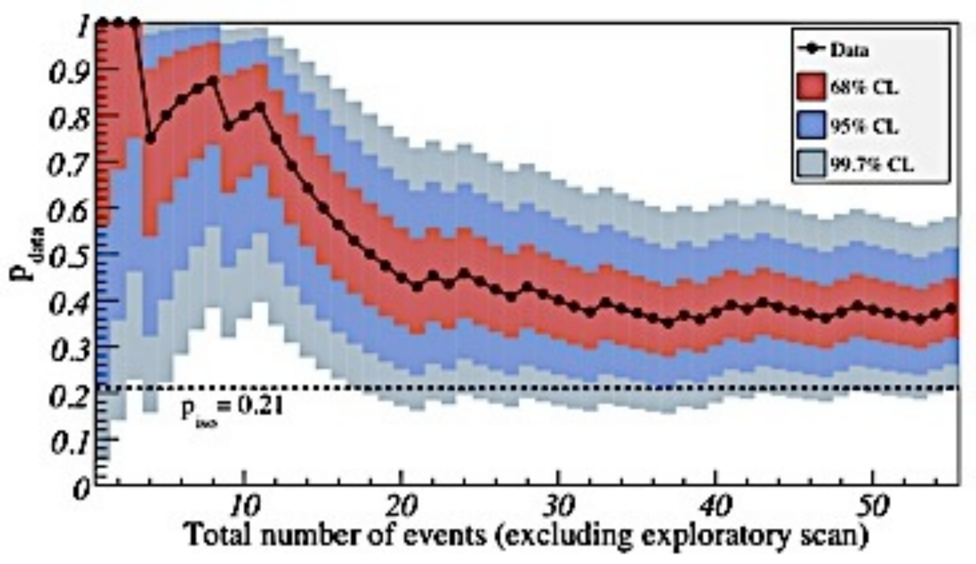


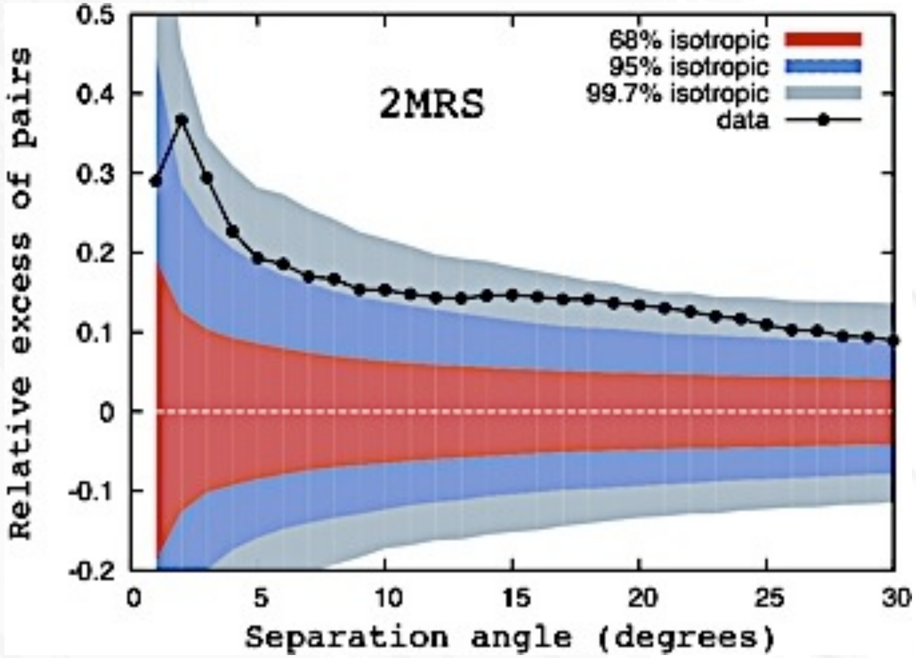
Observed events consistent with constraints on correlated sources for heavy primaries !

Ptitsyna, Troitsky., arXiv:0808.0367

Ultra-High Energy Cosmic Ray Sky Distribution

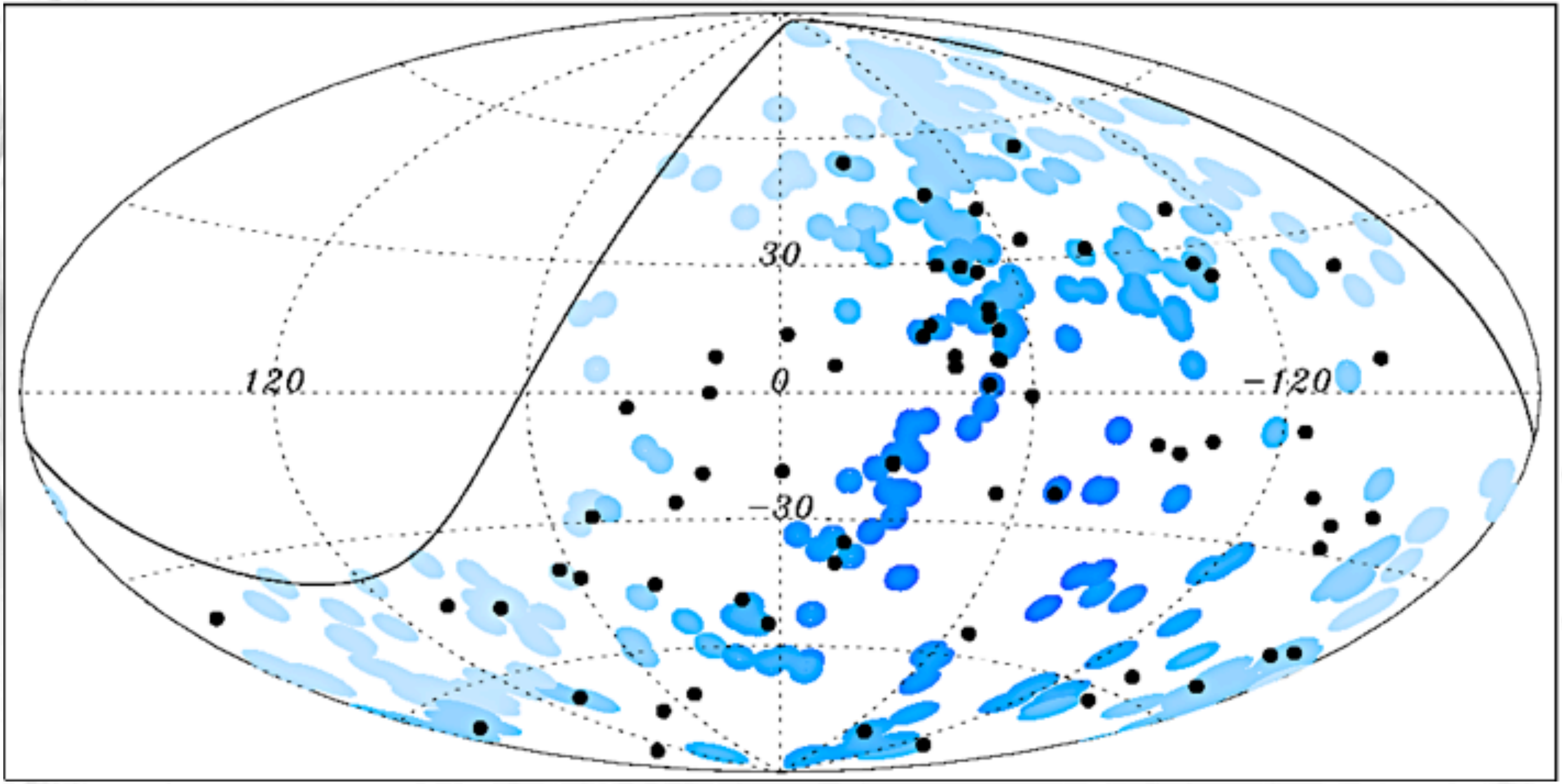
Pierre Auger Observatory update on correlations with nearby extragalactic matter: Pierre Auger Collaboration, arXiv:1009.1855





The case for anisotropy does not seem to have strengthened with more data: Fraction of events above 55 EeV correlating with the Veron Cetty Catalog has came down from 69+11-13% to 38+7-6% with 21% expected for isotropy. Excess of correlation also seen with 2MRS catalog at 95% CL.

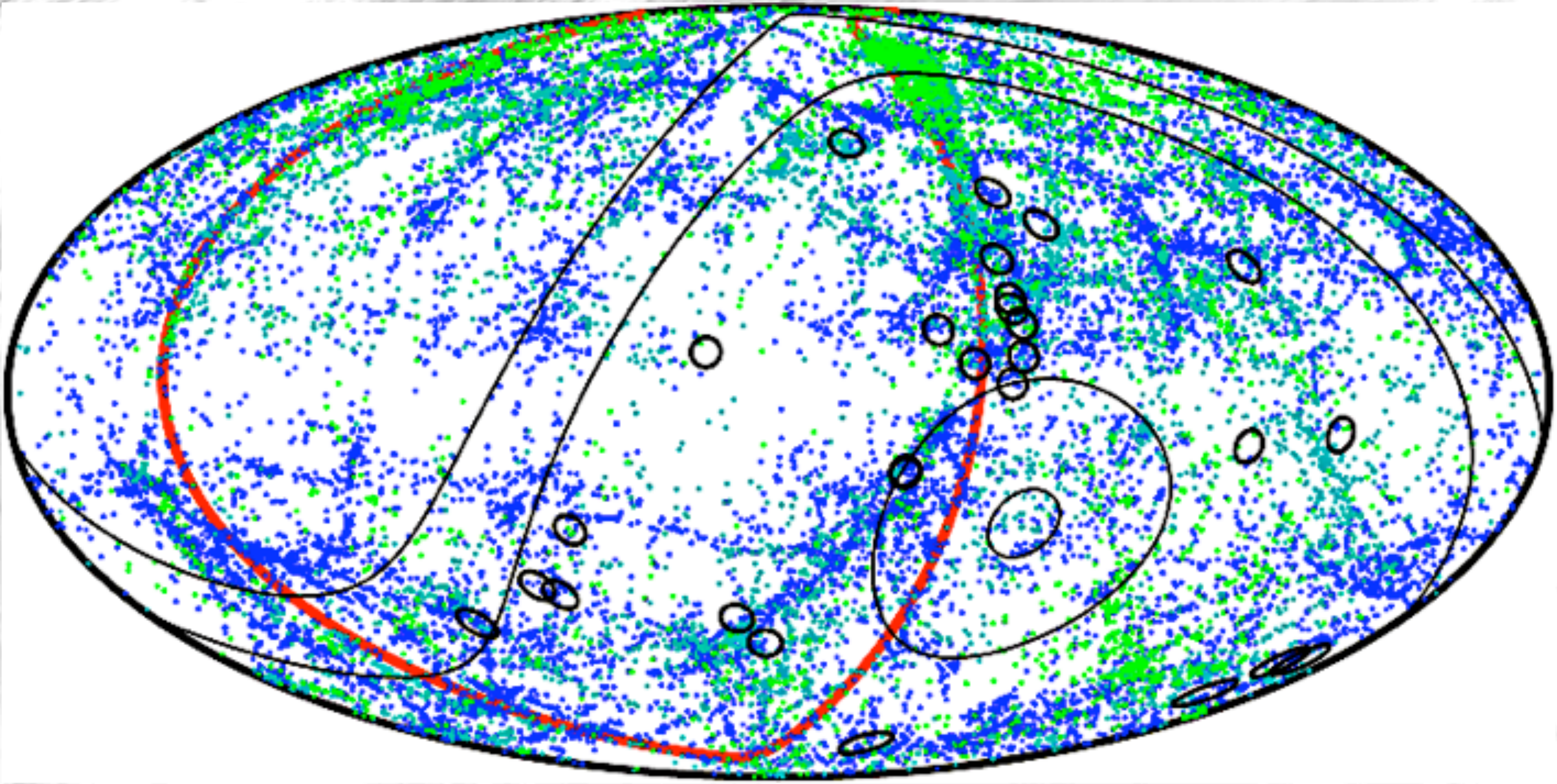
Auger sees Correlations with AGNs !



Blue 3.1 deg. circles = 318 AGNs from the Veron Cetty catalogue within 75 Mpc (exposure weighted color); black dots = 69 events above 55 EeV.

29 events correlated within 3.1°, 14.5 expected for isotropy

Pierre Auger Collaboration, arXiv:1009.1855



Points = galaxies with z < 0.015

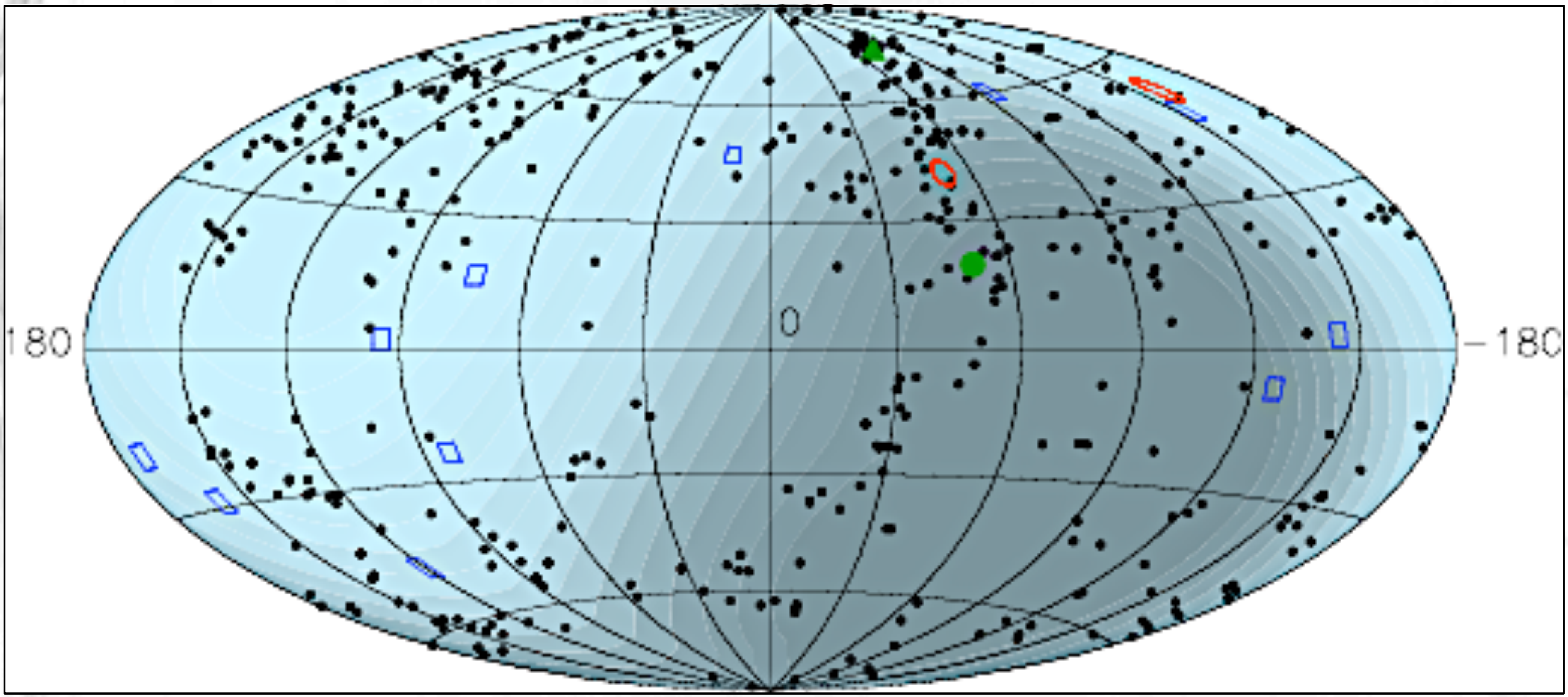
Black circles = Auger events above 60 EeV.

Black lines = equal exposure contours

red line= supergalactic plane

Lipari, arXiv:0808.0417

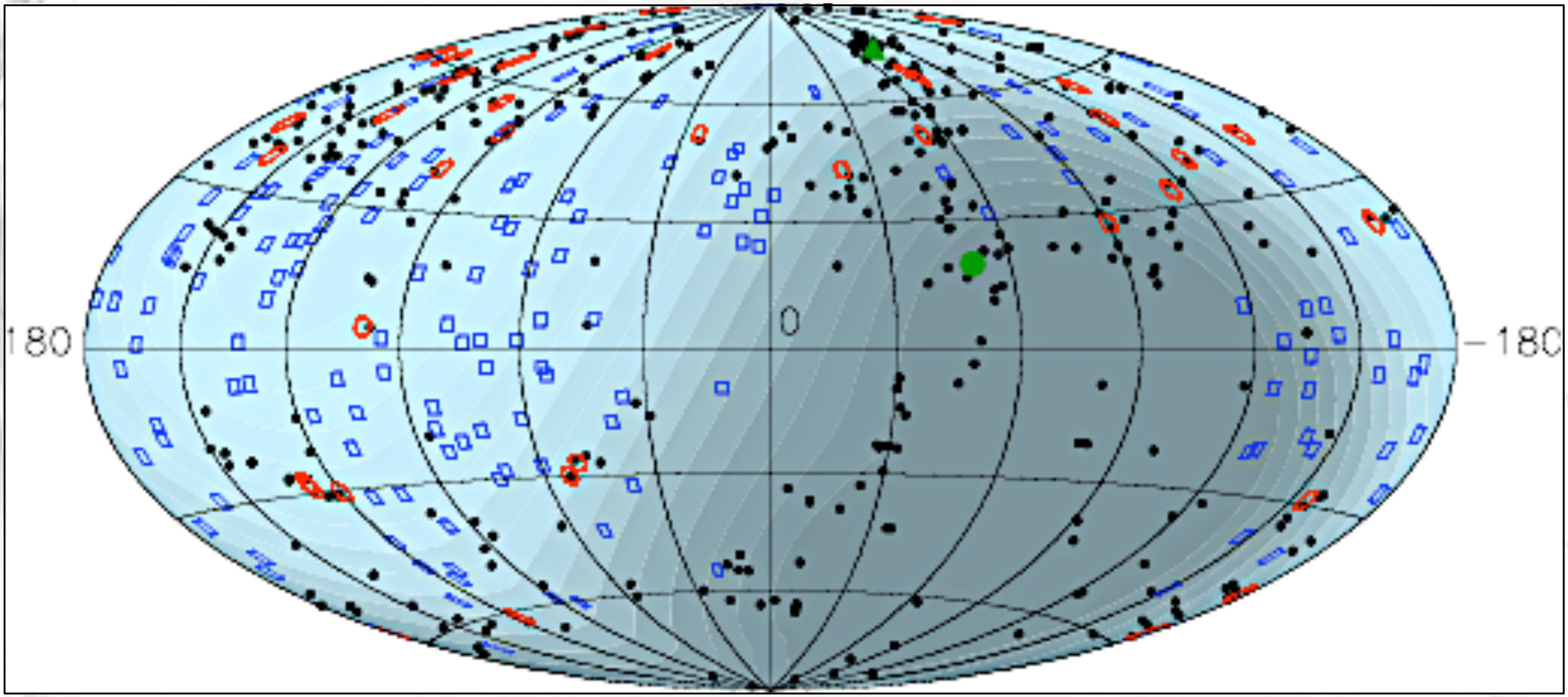
But HiRes sees no Correlations !



Black dots = 457 AGNs + 14 QSOs from the Veron Cetty catalogue for z < 0.018
red circles = 2 correlated events above 56 EeV within 3.1°,
blue squares = 11 uncorrelated events

HiRes Collaboration, Astropart.Phys. 30 (2008) 175

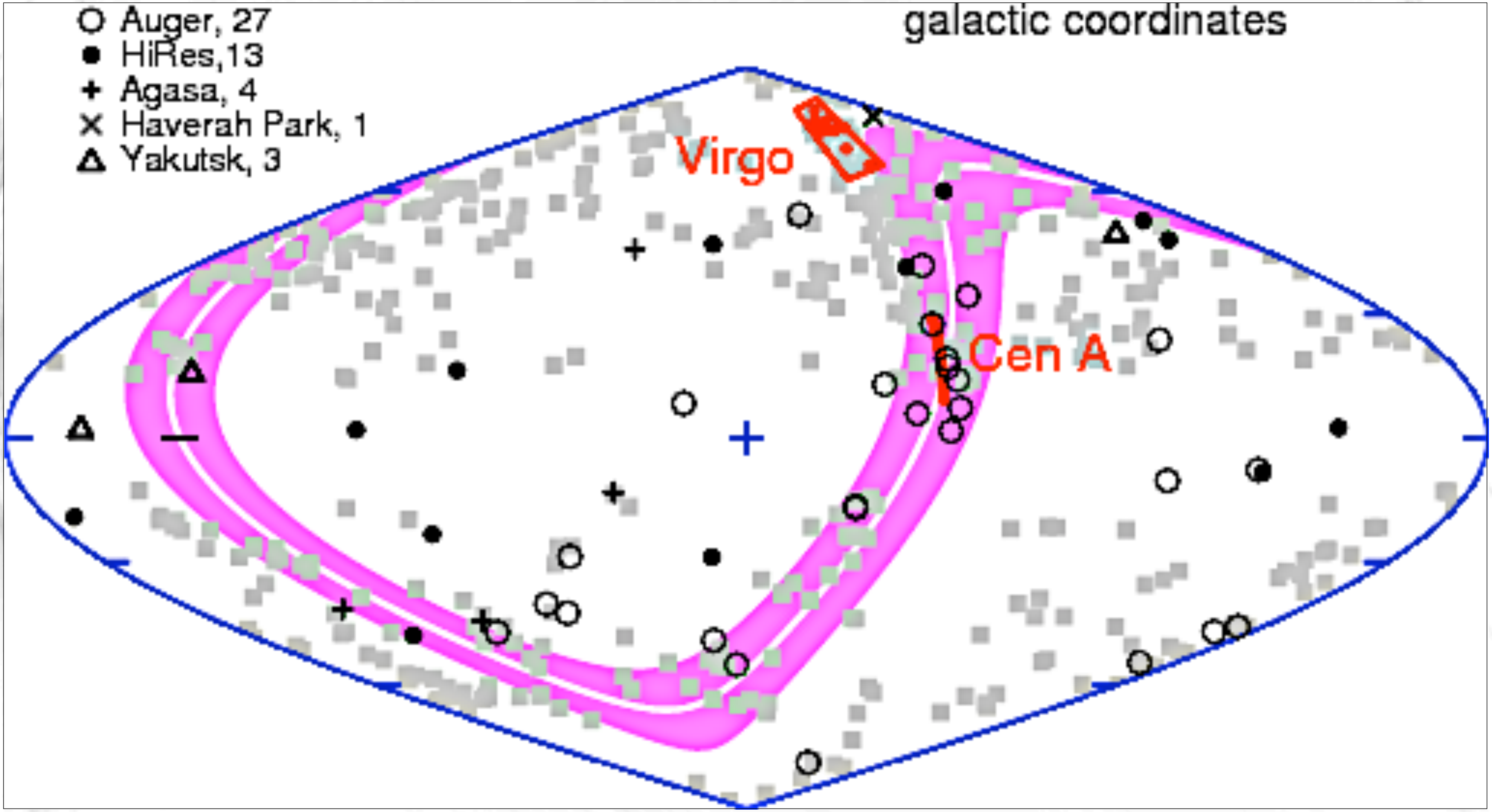
But HiRes sees no Correlations !



Black dots = 389 AGNs + 14 QSOs from the Veron Cetty catalogue for z < 0.016
red circles = 36 correlated events above 15.8 EeV within 2.0°,
blue squares = 162 uncorrelated events

HiRes Collaboration, Astropart.Phys. 30 (2008) 175

Correlation with supergalactic plane



Correlation with supergalactic plane within 10° (15°) is improved from 2.0 (2.4) sigma to 3.6 (3.2) sigma when definition relates to structure within 70 Mpc.

Stanev, arXiv:0805.1746

Further Curiosities in the Sky Distributions

too few events from Virgo cluster, see
Gorbunov et al., JETP Lett. 87 (2007) 461

too many events from Centaurus A, e.g. Moskalenko et al., arXiv:0805.1260; Rachen, arXiv:0808.0348.

The AGNs with which Auger events correlate are not thought to be strong enough, see Moskalenko et al., arXiv:0805.1260; Zaw, Farrar, Greene, arXiv:0806.3470 (the latter arguing for flares)

According to Gureev and Troitsky, arXiv:0808.0481, the correlation of Auger events with AGNs is stronger when nearest neighbor sources only are counted, than when all AGN within given off-set are counted. According to them, this reveals individual sources rather than the population.

Some general estimates for sources

Accelerating particles of charge eZ to energy E_{max} requires induction $\varepsilon > E_{max}/eZ$. With $Z_0 \sim 100\Omega$ the vacuum impedance, this requires dissipation of minimum power of

$$L_{\min} \sim \frac{\epsilon^2}{Z_0} \simeq 10^{45}\, Z^{-2} \left(\frac{E_{\max}}{10^{20}\,\mathrm{eV}}\right)^2 \mathrm{erg\,s^{-1}}$$

This „Poynting" luminosity can also be obtained from $L_{min} \sim (BR)^2$ where BR is given by the „Hillas criterium":

$$BR > 3 \times 10^{17}\,\Gamma^{-1} \left(\frac{E_{\max}}{10^{20}\,\mathrm{eV}}\right) \mathrm{Gauss\,cm}$$

where Γ is a possible beaming factor.
If most of this goes into electromagnetic channel, only AGNs and maybe gamma-ray bursts could be consistent with this.

In arXiv:1003.2500 Hardcastle estimates a corresponding lower limit on the radio luminosity:

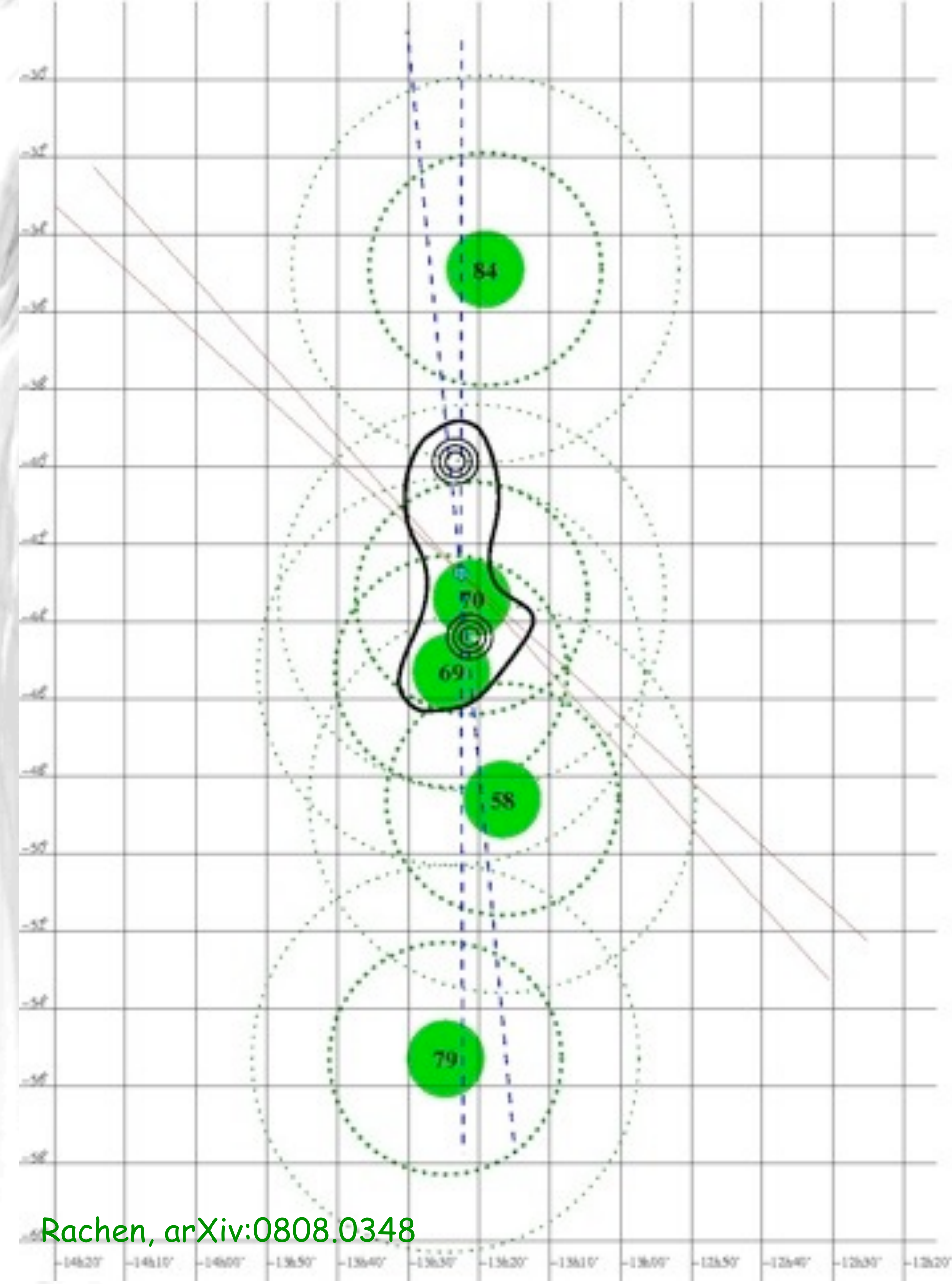
$$L_{108\,\mathrm{MHz}} > 2 \times 10^{24}\,\epsilon \left(\frac{E/Z}{10^{20}\,\mathrm{eV}}\right)^{7/2} \left(\frac{r_{\mathrm{lobe}}}{100\,\mathrm{kpc}}\right)^{-1/2} \mathrm{W\,Hz^{-1}}$$

for an E^{-2} electron spectrum
with ε = energy in electrons / energy in magnetic field

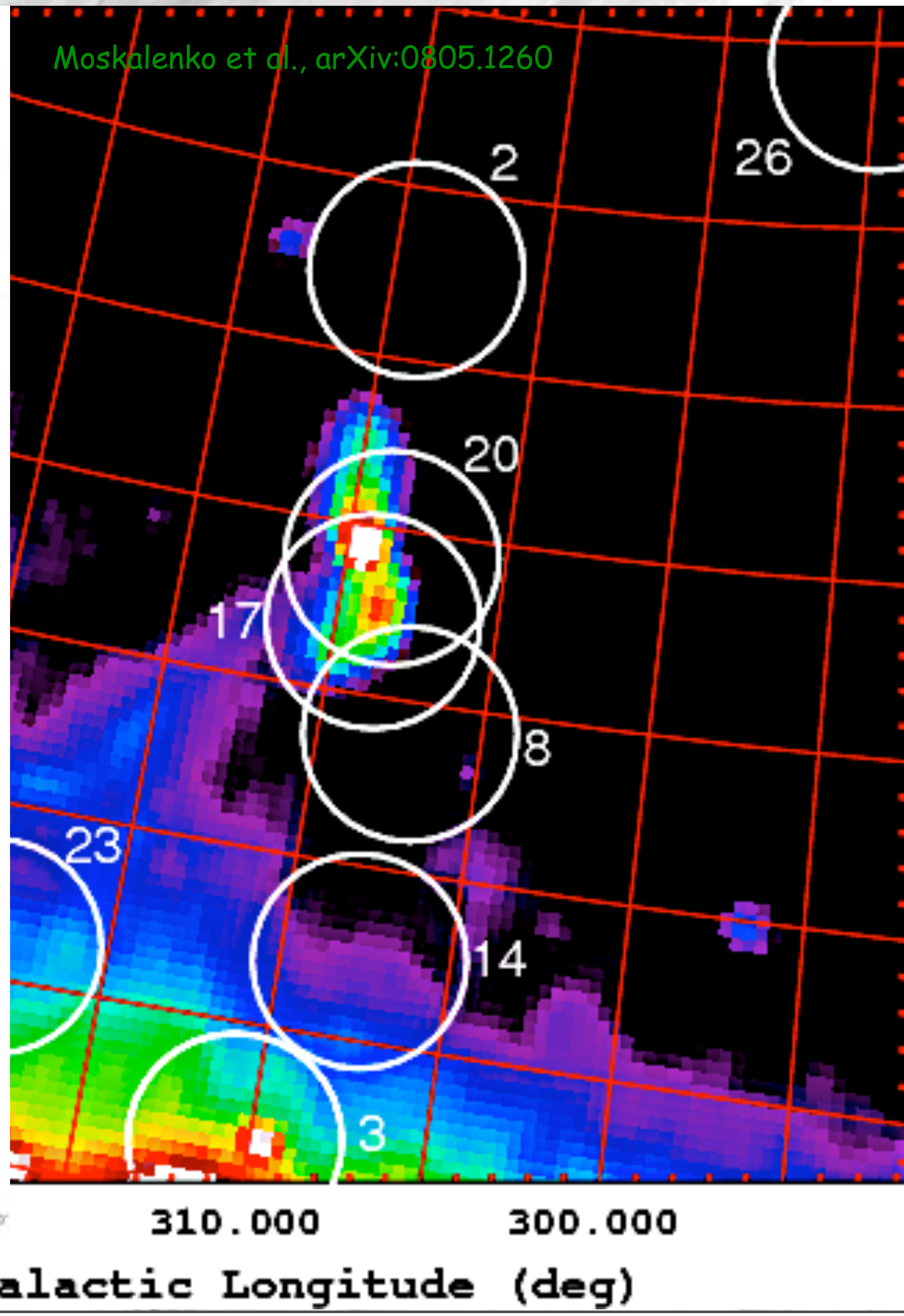
He concludes: if protons, then very few sources which should be known and spectrum should cut off steeply at observed highest energies

If heavier nuclei then there are many radio galaxy sources but only Cen A may be identifiable

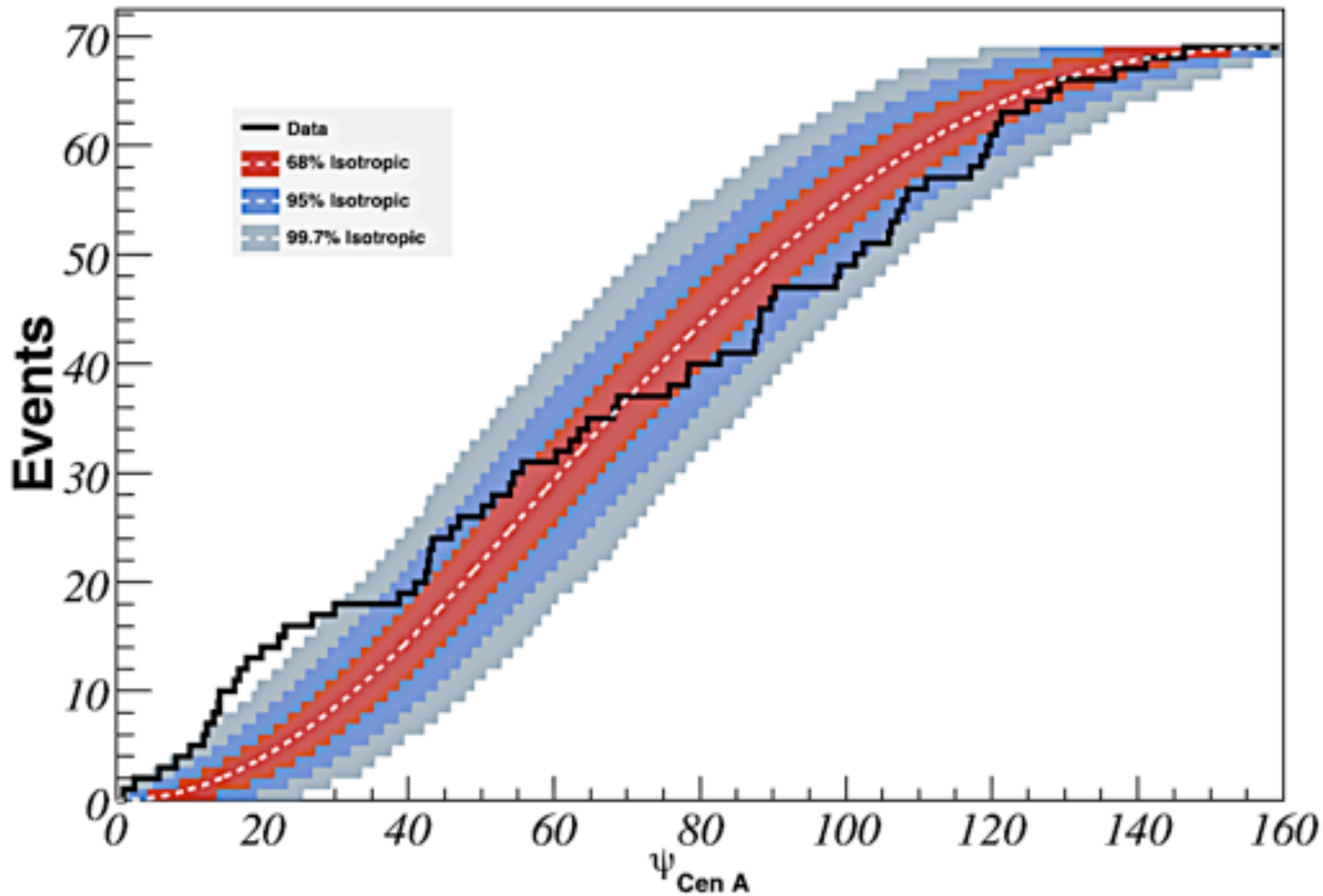
Centaurus A



84
70
69
58
79
Rachen, arXiv:0808.0348



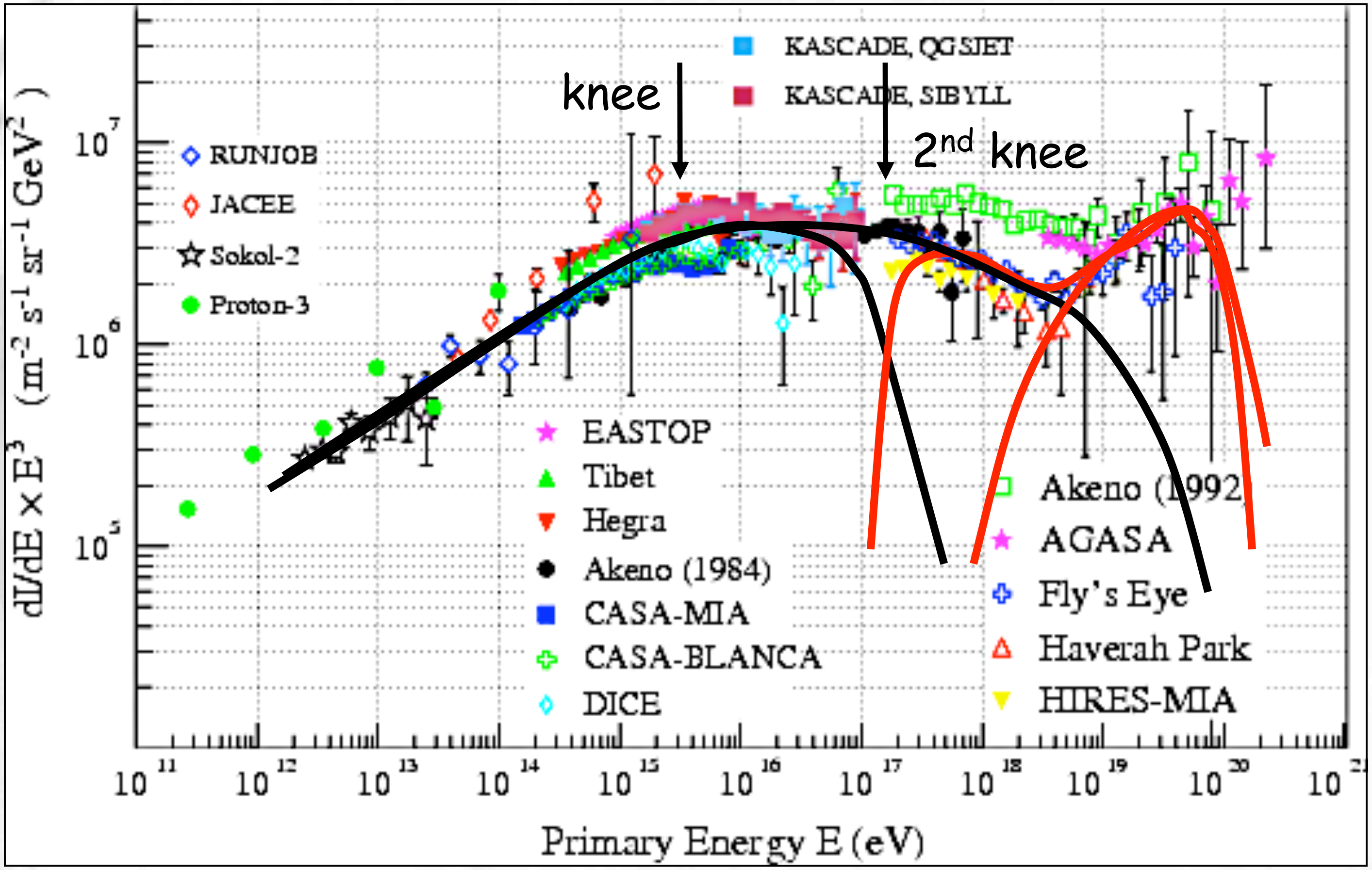
Moskalenko et al., arXiv:0805.1260
2
26
20
17
8
23
14
3
310.000
300.000
Galactic Longitude (deg)



Pierre Auger sees a clear excess in the direction of Centaurus A.

Pierre Auger Collaboration, arXiv:1009.1855

Chemical Composition, Nature of the Ankle



"Conventional Scenario":

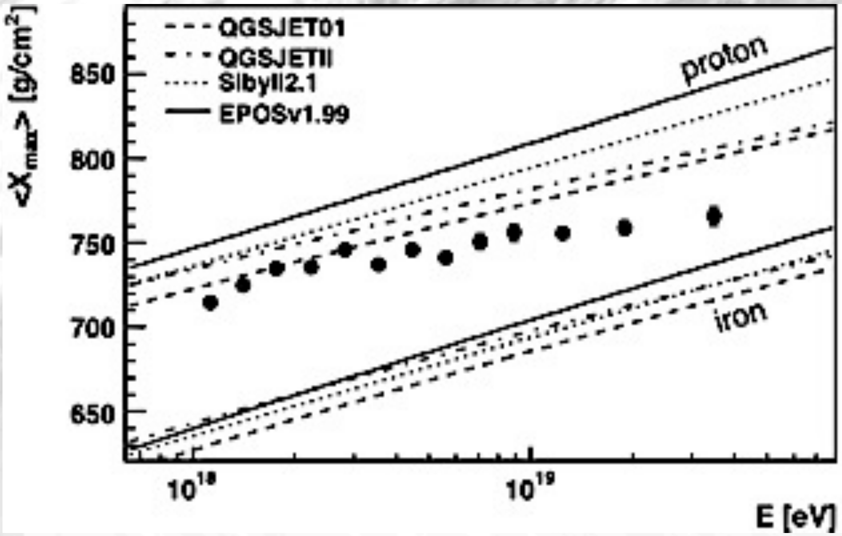
The ankle at ~5x10^{18} eV is a cross-over from a heavy Galactic to a light extragalactic component.

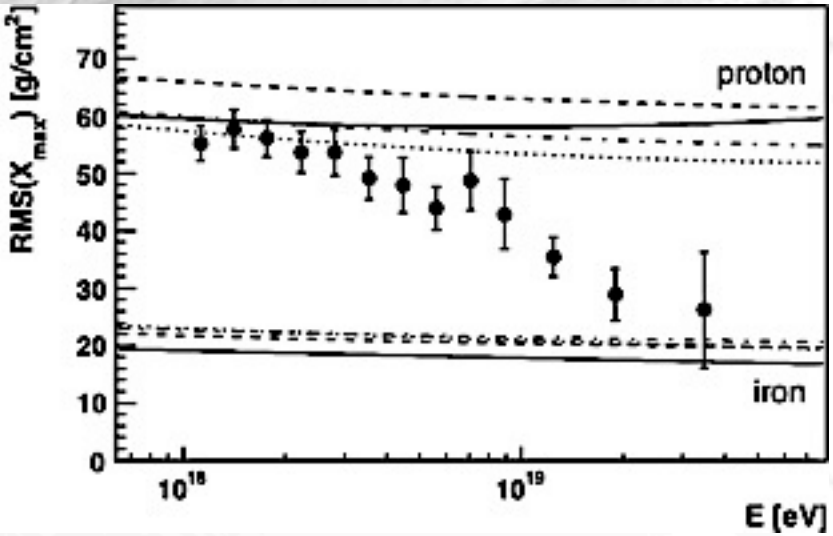
Scenario of Berezinsky et al.:

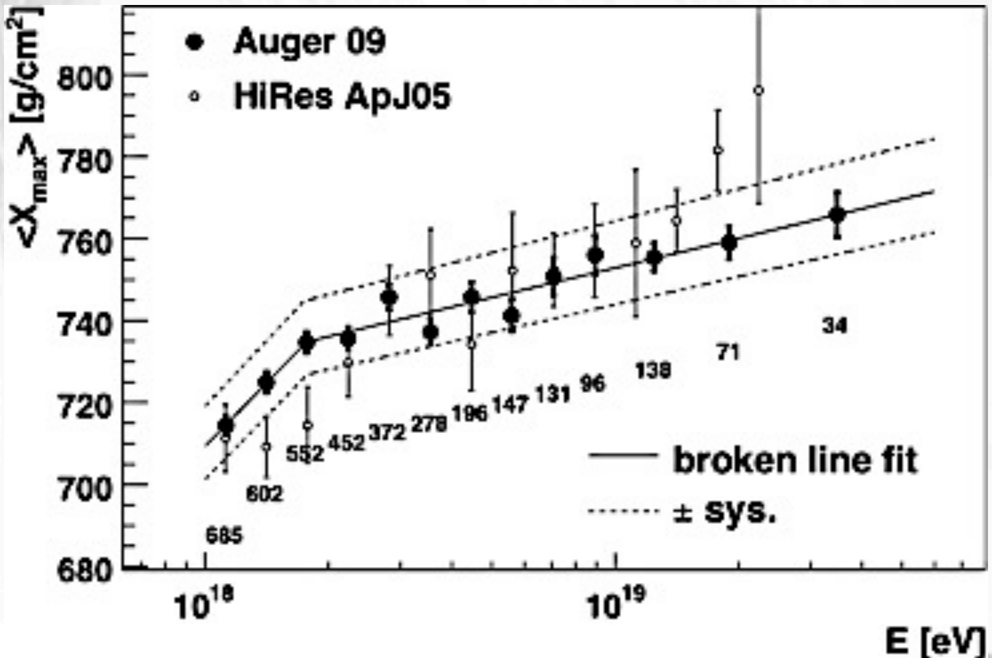
Galactic cosmic rays level out at the 2nd knee at ~4x10^{17} eV where dominated by heavy nuclei.

The ankle at ~5x10^{18} eV is due to pair production of extragalactic protons on the CMB. Requires >85% protons at the ankle.

There may be a significant heavy component at the highest energies:



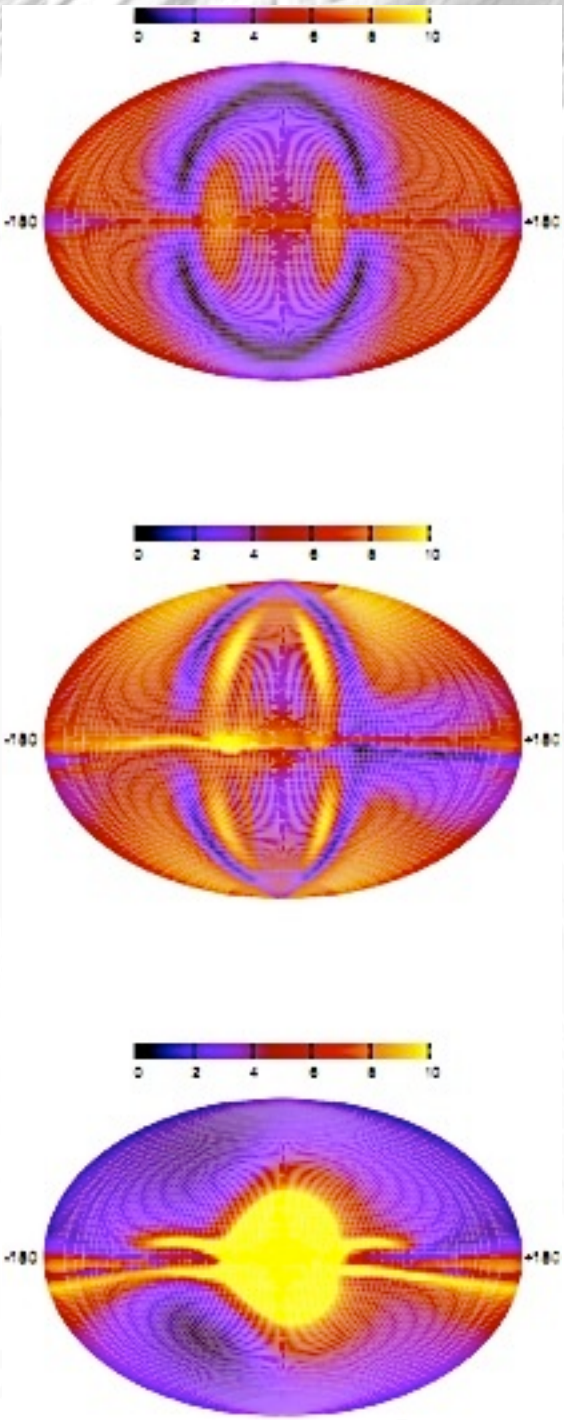




Auger data on composition seem to point to a quite heavy composition at the highest energies, whereas HiRes data seem consistent with a light composition.

Pierre Auger Collaboration,
Phys.Rev.Lett., 104 (2010) 091101

HiRes Collaboration,
Phys.Rev.Lett. 104 (2010) 161101



Consequences for Galactic Deflection

Deflection in **galactic magnetic field** is rather model dependent, here for E/Z=4 10^{19} eV for Models of

Tinyakov, Tkachev (top)

Harrari, Mollerach, Roulet (middle)

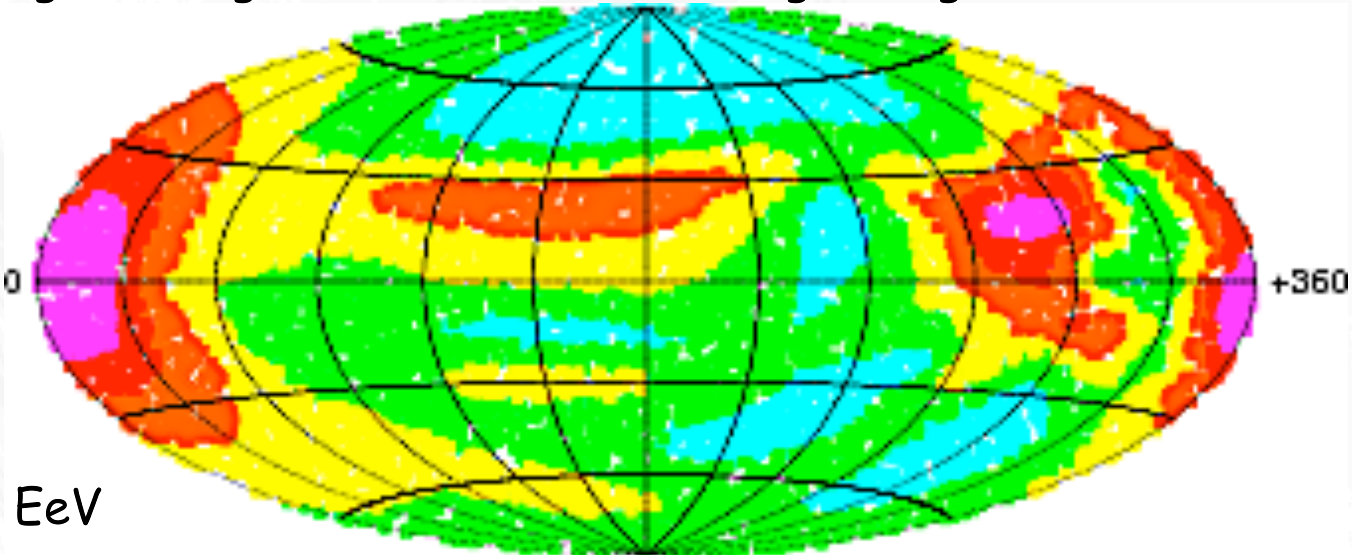
Prouza, Smida (bottom)

Deflection in **extragalactic fields** is even more uncertain

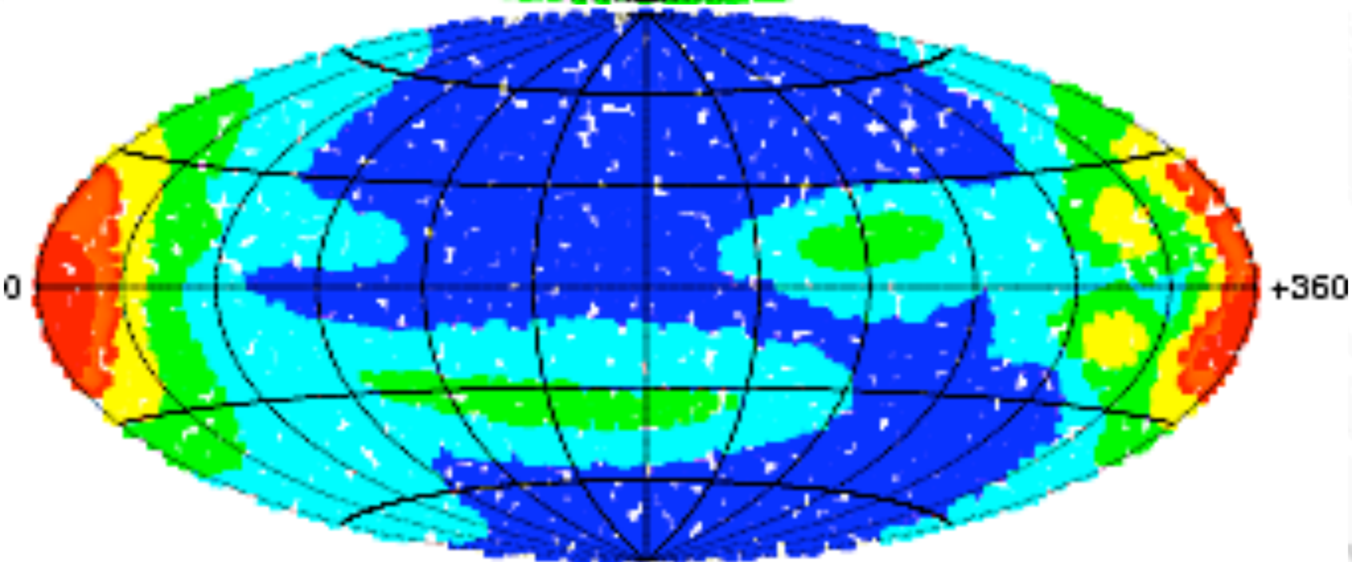
Kachelriess, Serpico, Teshima, Astropart. Phys. 26 (2006) 378

Deflection of iron in galactic magnetic field model of Prouza&Smida

Angular range between 0 and 100 degrees, galactic coordinates



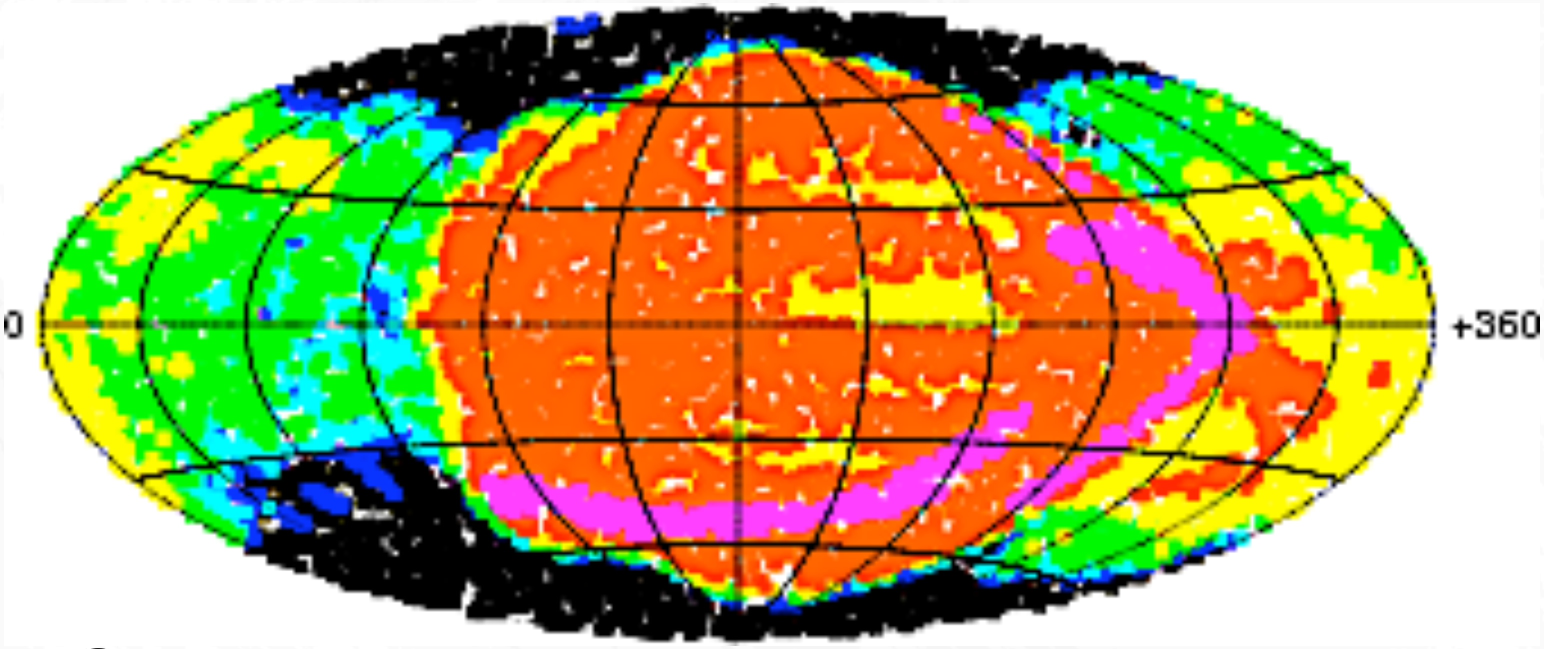
E=60 EeV



E=140 EeV

Giacinti, Kachelriess, Semikoz, Sigl, JCAP 1008 (2010) 036

Bachtracking of iron in galactic magnetic field model of Prouza&Smida



E=60 EeV

Giacinti, Kachelriess, Semikoz, Sigl, JCAP 1008 (2010) 036

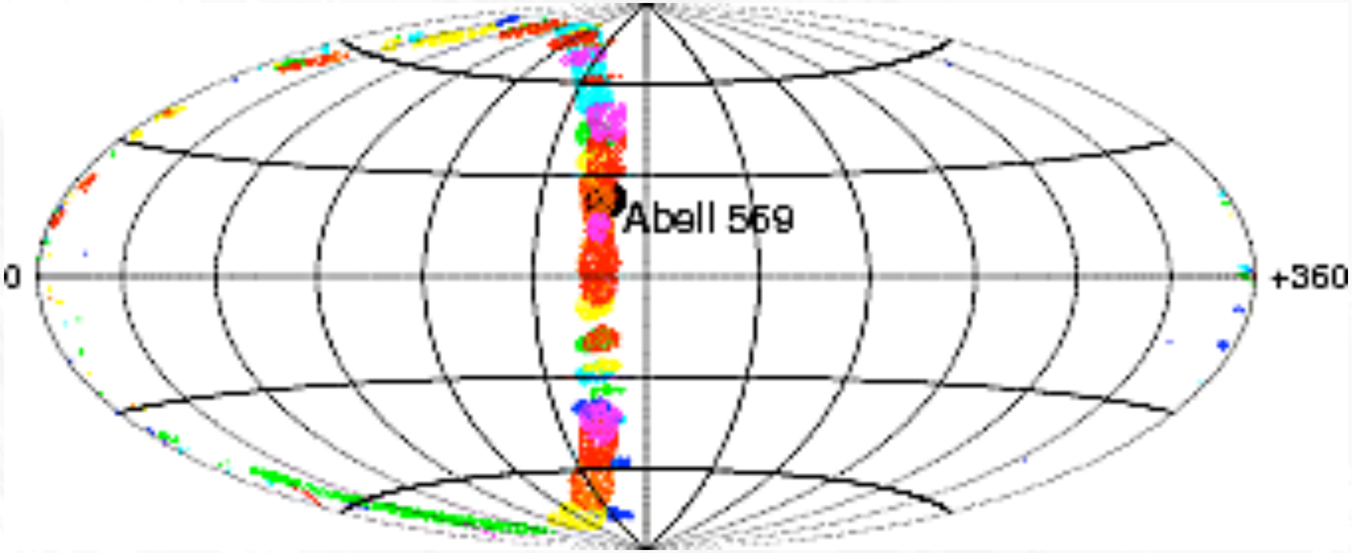
Density range between 10^{-3} and $10^{0.5}$, galactic coordinates

Highly anisotropic picture

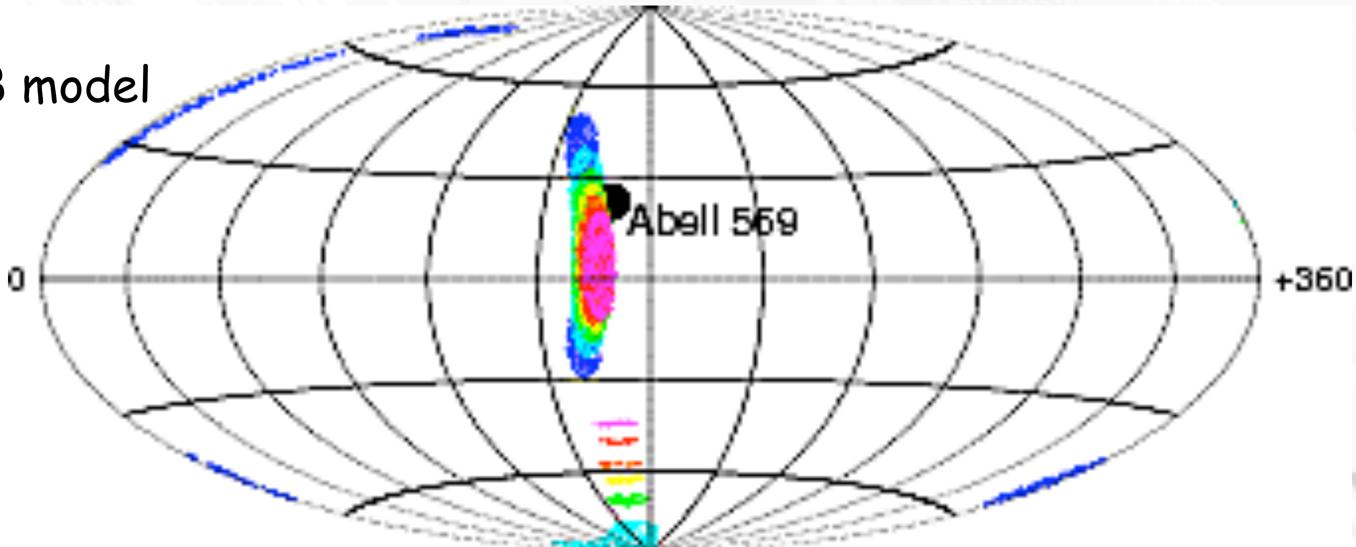
Empty backtracked regions are invisible from within the Galaxy !

“Iron Image” of galaxy cluster Abell0569 in two galactic field models

Energy range from 60 to 140 EeV

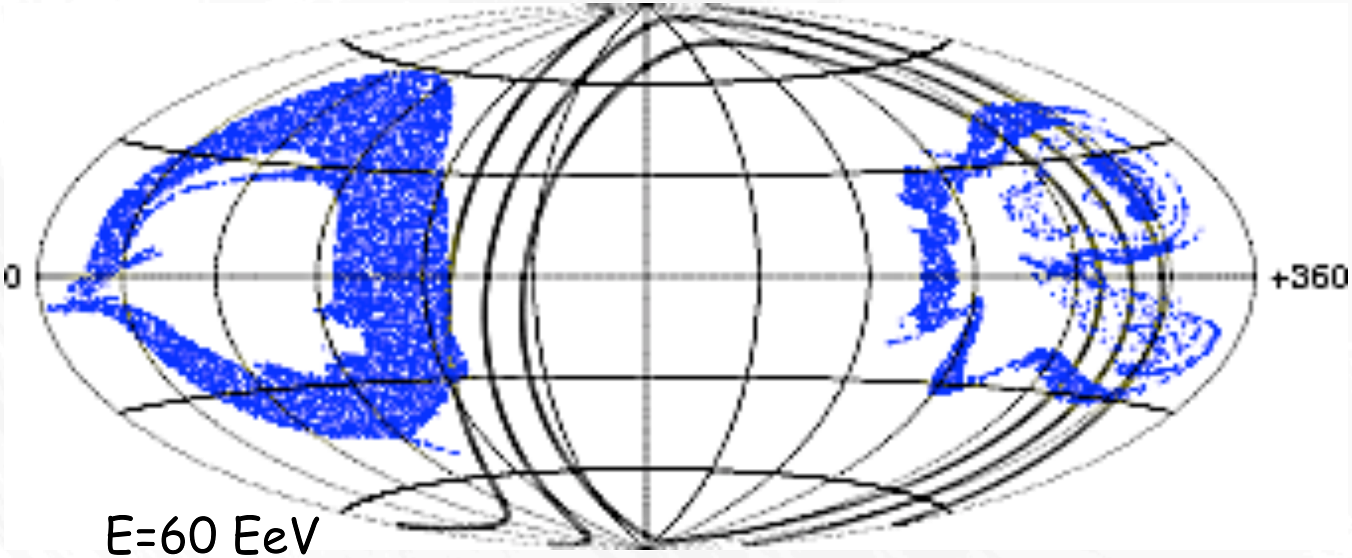


Sun08 model

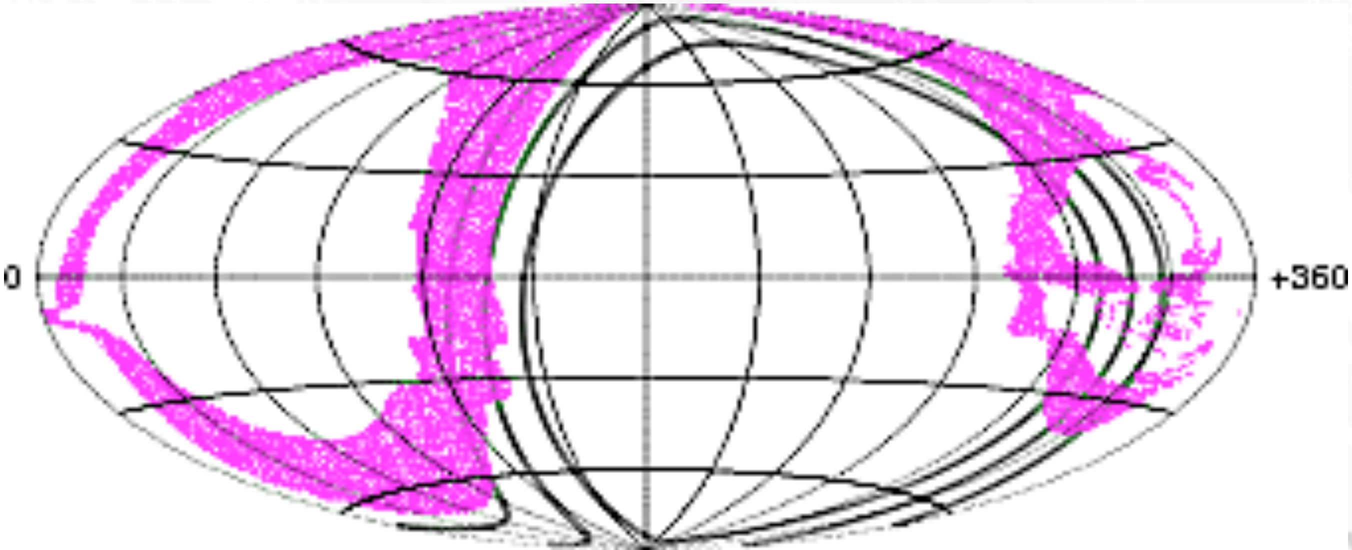


Giacinti, Kachelriess, Semikoz, Sigl, JCAP 1008 (2010) 036

"Iron image" of supergalactic plane in galactic magnetic field model of Prouza&Smida



E=60 EeV



E=140 EeV

Giacinti, Kachelriess, Semikoz, Sigl, JCAP 1008 (2010) 036

“Conundrum”:

If deflection is small and sources follow the local large scale structure then

a) primaries should be protons to avoid too much deflection in galactic field

b) but air shower measurements by Pierre Auger (but not HiRes) indicate mixed or heavy composition

c) Theory of AGN acceleration seem to necessitate heavier nuclei to reach observed energy

Extragalactic Ultra-High Energy Cosmic Ray Propagation and Magnetic Fields

Cosmic rays above ~10^{19} eV are probably extragalactic and may be deflected mostly by extragalactic fields B_{XG} rather than by galactic fields.

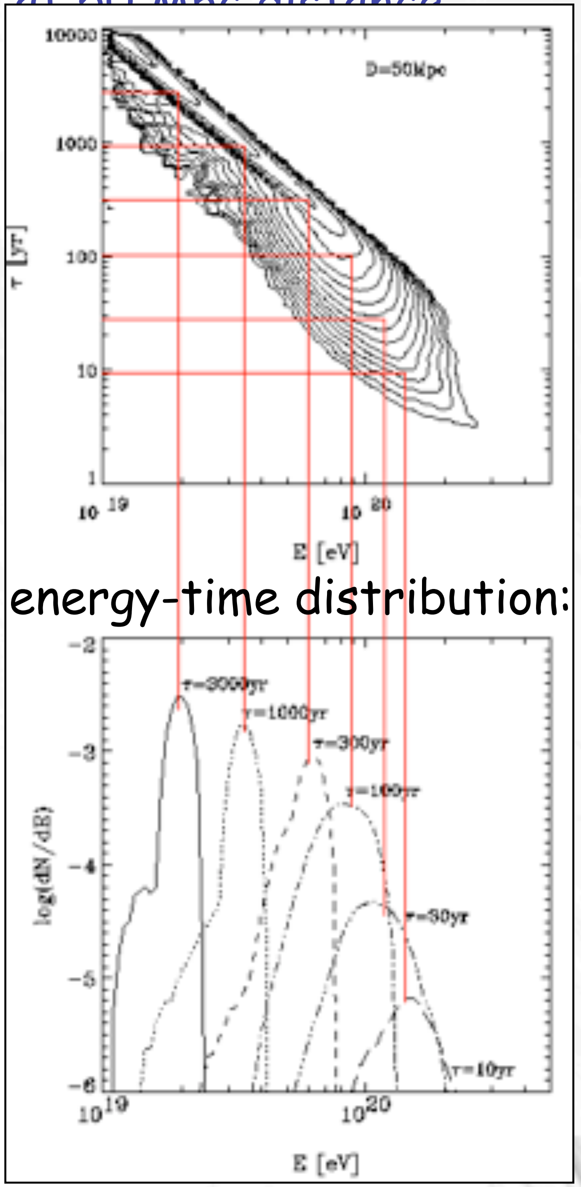
However, very little is known about about B_{XG}: It could be as small as 10^{-20} G (primordial seeds, Biermann battery) or up to fractions of micro Gauss if concentrated in clusters and filaments (equipartition with plasma).

Transition from rectilinear to diffusive propagation over distance d in a field of strength B and coherence length λ_c at:

$$E_c \sim 1.2 \times 10^{19} \left(\frac{Z}{26}\right) \left(\frac{d}{\mathrm{Mpc}}\right)^{1/2} \left(\frac{B_{\mathrm{rms}}}{\mathrm{nG}}\right) \left(\frac{\lambda_c}{\mathrm{Mpc}}\right)^{1/2} \mathrm{eV}$$

Example: Magnetic field of 10^{-10} Gauss,
coherence scale 1 Mpc,
burst source at 50 Mpc distance

time delay

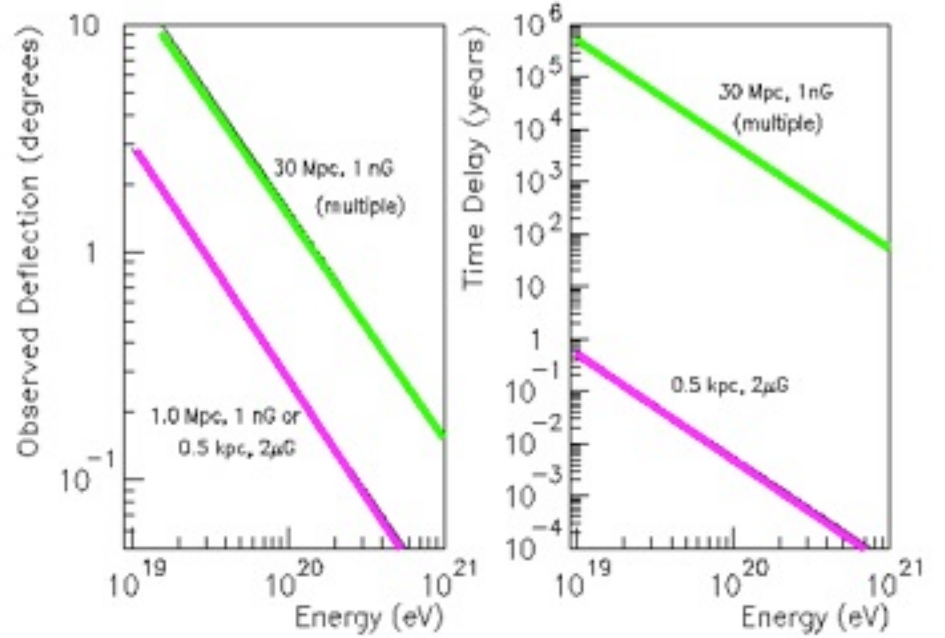


cuts through energy-time distribution:

differential spectrum

Lemoine, Sigl

Typical numbers:



Δθ for a vertical shower:

	10 EeV	100 EeV
Array alone	2°	<1°
Hybrid	0.25°	0.20°

Transition rectilinear-diffusive regime

Neglect energy losses for simplicity.

Time delay over distance d in a field B_{rms} of coherence length λ_c for small deflection:

$$\tau(E,d) \simeq \frac{d\theta(E,d)^2}{4} \simeq 1.5\times 10^3\, Z^2 \left(\frac{E}{10^{20}\,\mathrm{eV}}\right)^{-2}\left(\frac{d}{10\,\mathrm{Mpc}}\right)^2\left(\frac{B_{\rm rms}}{10^{-9}\,\mathrm{G}}\right)^2\left(\frac{\lambda_c}{\mathrm{Mpc}}\right)\,\mathrm{yr}$$

This becomes comparable to distance d at energy E_c:

$$E_c \sim 4.7\times 10^{19}\, Z\left(\frac{d}{10\,\mathrm{Mpc}}\right)^{1/2}\left(\frac{B_{\rm rms}}{10^{-7}\,\mathrm{G}}\right)\left(\frac{\lambda_c}{\mathrm{Mpc}}\right)^{1/2}\,\mathrm{eV}$$

In the rectilinear regime for total differential power Q(E) injected inside d, the differential flux reads

$$j(E) = \frac{Q(E)}{(4\pi d)^2}$$

In the diffusive regime characterized by a diffusion constant D(E), particles are confined during a time scale

$$\tau(E,d) \simeq \frac{d^2}{D(E)}$$

which leads to the flux

$$j(E) \simeq \frac{Q(E)\tau(E)}{(4\pi)^2 d^3} = \frac{Q(E)}{(4\pi)^2 d D(E)}$$

For a given power spectrum B(k) of the magnetic field an often used (very approximate) estimate of the diffusion coefficient is

$$D(E) \simeq \frac{r_g(E)}{3} \frac{B_{\rm rms}}{\int_{1/r_g(E)}^{\infty} dk k^2 \langle B^2(k)\rangle},$$

where $B_{rms}^2 = \int_0^\infty dk k^2 \langle B^2(k) \rangle$, and the gyroradius is

$$r_g(E) \simeq \frac{E}{ZeB_{\rm rms}} \simeq 110\, Z^{-1} \left(\frac{E}{10^{20}\,{\rm eV}}\right) \left(\frac{B_{\rm rms}}{10^{-6}\,{\rm G}}\right)^{-1} {\rm kpc}$$

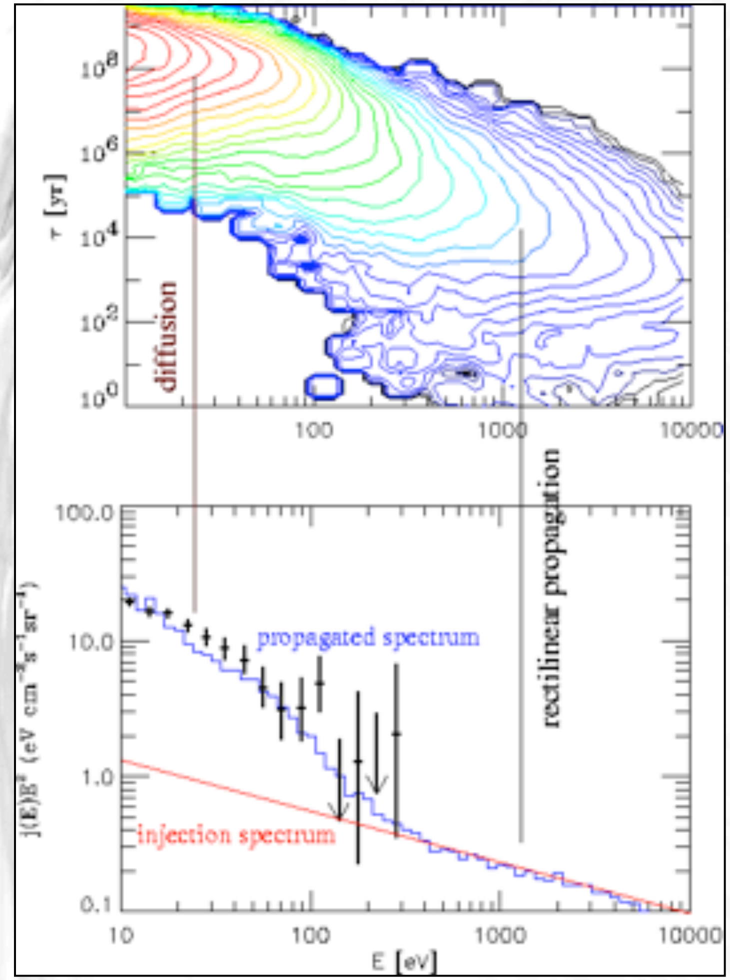
IF $E \ll E_c$ and IF energy losses can be approximated as continuous, $dE/dt=-b(E)$ (this is not the case for pion production), the local cosmic ray density $n(E,\mathbf{r})$ obeys the diffusion equation

$$\partial_t n(E,\mathbf{r}) + \partial_E \left[b(E)n(E,\mathbf{r})\right] - \nabla \cdot \left[D(E,\mathbf{r})\nabla n(E,\mathbf{r})\right] = q(E,\mathbf{r})$$

Where now $q(E,\mathbf{r})$ is the differential injection rate per volume, $Q(E)=\int d^3\mathbf{r}\, q(E,\mathbf{r})$. Analytical solutions exist (Syrovatskii), but the necessary assumptions are in general too restrictive for ultra-high energy cosmic rays.

Monte Carlo codes are therefore in general indispensable.

Transition rectilinear-diffusive regime: Summary



$$\tau(E) \propto d\theta^2 \propto \frac{d^2}{E^2}$$ in rectilinear regime

$$\tau(E,d) \simeq \frac{d^2}{D(E)}$$ in diffusive regime

$$j(E) \propto \frac{Q(E)}{d^2}$$ in rectilinear regime

$$j(E) \propto \frac{Q(E)\tau(E)}{d^3} \propto \frac{Q(E)}{dD(E)}$$ in diffusive regime

Simulated example: Continuous source distribution following Gaussian profile; $B=3\times10^{-7}$ G, d=10 Mpc, λ_c=1 Mpc.

Transition at energy $$E_c \sim 4.7\times10^{19}\,Z\left(\frac{d}{10\,\mathrm{Mpc}}\right)^{1/2}\left(\frac{B_{\rm rms}}{10^{-7}\,\mathrm{G}}\right)\left(\frac{\lambda_c}{\mathrm{Mpc}}\right)^{1/2}\,\mathrm{eV}$$

In the transition regime Monte Carlo codes are in general indispensable.

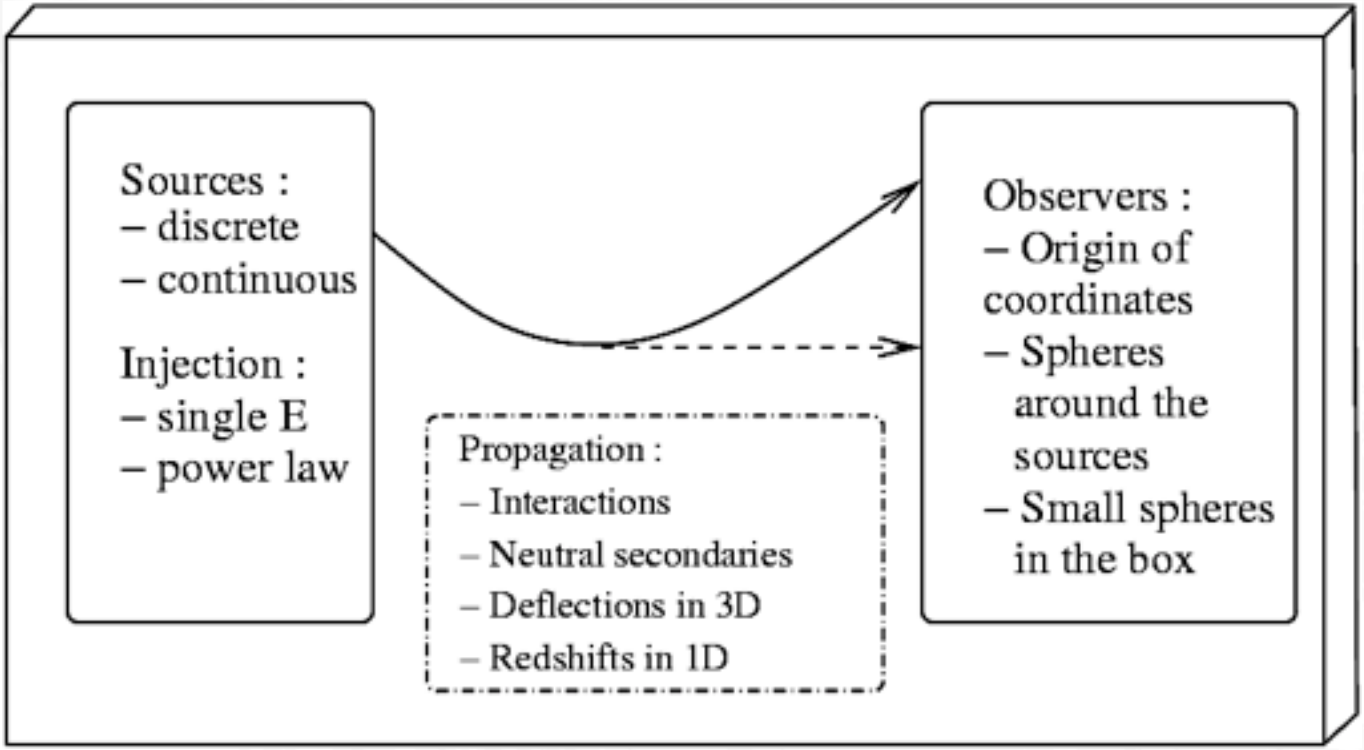
Principle of deflection Monte Carlo code

A particle is registered every time a trajectory crosses the sphere around the observer. This version to be applied for individual source/magnetic field realizations and inhomogeneous structures.

Main Drawback: CPU-intensive if deflections are considerable because most trajectories are “lost”. But inevitable for accurate simulations in highly structured environments without symmetries.

Simulating Propagation of Ultrahigh Energy Cosmic Rays, Gamma-Rays and Neutrinos with CRPropa

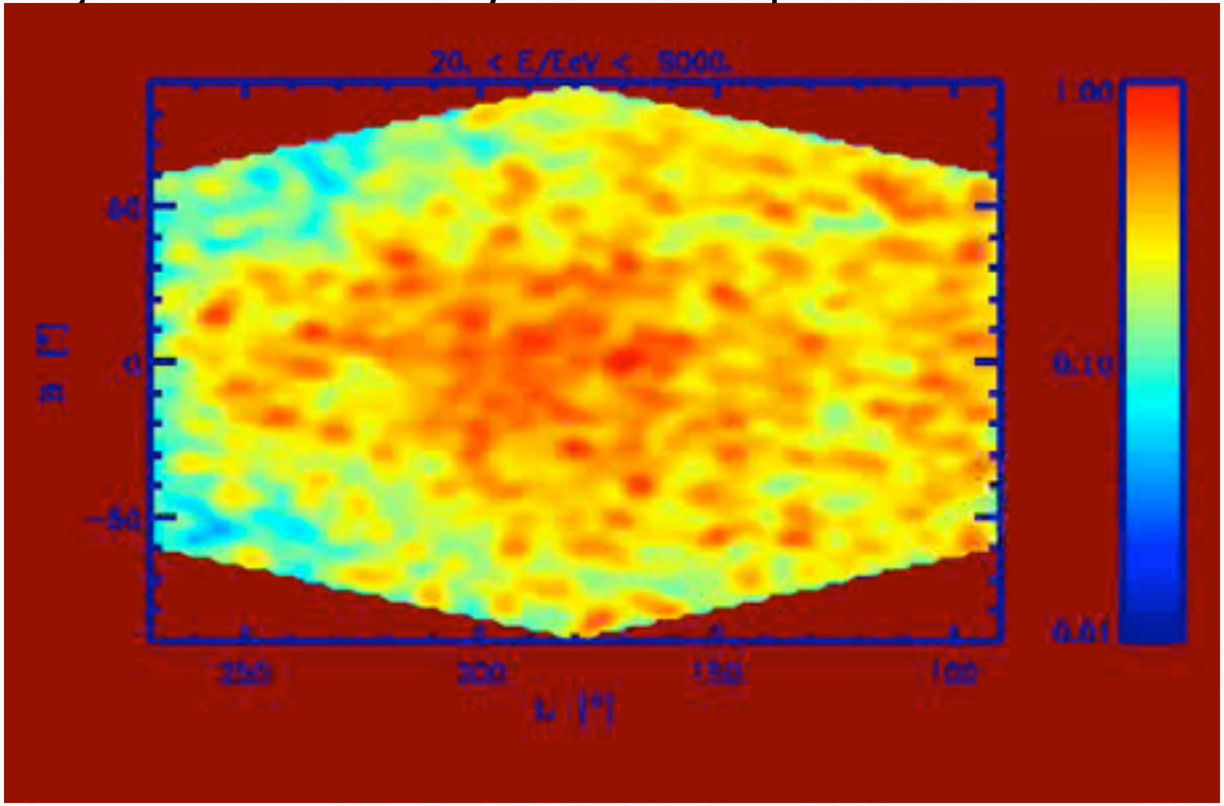
CRPropa is a public code for UHE cosmic rays, neutrinos and γ-rays being extended to heavy nuclei and hadronic interactions



Eric Armengaud, Tristan Beau, Günter Sigl, Francesco Miniati,
Astropart.Phys.28 (2007) 463.
http://apcauger.in2p3.fr/CRPropa/index.php
Now including: Jörg Kulbartz, Luca Maccione, Ricard Tomas, Mariam Tortola,
Nils Nierstenhoefer, Karl-Heinz Kampert, ...

Effects of a single source: Numerical simulations

A source at 3.4 Mpc distance injecting protons with spectrum $E^{-2.4}$ up to 10^{22} eV
A uniform Kolmogorov magnetic field, $\langle B^2(k)\rangle \sim k^{-11/3}$, of rms strength 0.3 μG,
and largest turbulent eddy size of 1 Mpc.



10^5 trajectories,
251 images between
20 and 300 EeV,
2.5° angular resolution

Isola, Lemoine, Sigl

Conclusions:

1.) Isotropy is inconsistent with only one source.

2.) Strong fields produce interesting lensing (clustering) effects.

The Universe is structured